

Collider Signatures of Higgs-portal Scalar Dark Matter

Huayong Han,¹ Jin Min Yang,^{2,3} Yang Zhang,² and Sibor Zheng¹

¹*Department of Physics, Chongqing University, Chongqing 401331, P. R. China*

²*Institute of Theoretical Physics, Academia Sinica, Beijing 100190, P. R. China*

³*Department of Physics, Tohoku University, Sendai 980-8578, Japan*

(Dated: Jan 2016)

In the simplest Higgs-portal scalar dark matter model, the dark matter mass has been restricted to be either near the resonant mass ($m_h/2$) or in a large-mass region by the direct detection at LHC Run 1 and LUX. While the large-mass region below roughly 3 TeV can be probed by the future Xenon1T experiment, most of the resonant mass region is beyond the scope of Xenon1T. In this paper, we study the direct detection of such scalar dark matter in the narrow resonant mass region at the 14 TeV LHC and the future 100 TeV hadron collider. We show the luminosities required for the 2σ exclusion and 5σ discovery.

I. INTRODUCTION

New physics beyond the Standard Model (SM) has drawn extensive attention since the discovery of the SM Higgs boson [1, 2]. While a few problems such as how to stabilize the Higgs mass against ultraviolet radiative corrections are tied to new physics of high mass scale, in this paper we instead focus on dark matter with a mass near the weak scale. In contrast to new physics which appears at a rather high mass scale, such a dark matter model has promising prospect for discovery at both astrophysical and particle collider experiments.

In particular, we are interested in the simplest Higgs-portal dark matter model, in which the dark matter communicates with SM particles via the Higgs scalar. Unlike the fermion dark matter setting, a scalar dark matter in the so-called Higgs-portal scalar dark matter model (HSDM) [3–7] still survives the latest data of direct detections at Xenon100 [8] and LUX [9], indirect detections at Fermi-LAT [10, 11], and Higgs invisible decay at the LHC Run 1 [12]. Detailed discussions about this model have been given in the literature ([13]–[42]). Fitting the experimental data indicates that the dark matter mass is either near the resonant mass region between 53 GeV and 62.5 GeV or in a large-mass region above 185 GeV.

While the large-mass region between 185 GeV and 3 TeV can be probed by the future Xenon1T [43], most of the resonant mass region is beyond the reach of this facility. In this paper, we discuss the collider signatures of the scalar dark matter in the HSDM model with a mass between 53 GeV and 62.5 GeV at the 14 TeV LHC and the future 100 TeV proton collider (FCC). We will show that similar to Circular Electron Positron Collider (CEPC) [44, 45], FCC will be a useful machine for searching dark matter in this narrow mass region. We will show that for FCC with a luminosity of 10 ab^{-1} the exclusion and discovery sensitivities reach to 57 GeV and 56 GeV respectively through the Vector Boson Fusion (VBF) channel, and 54.8 GeV and 53.9 GeV respectively via the mono- Z channel. It indicates that FCC with 10 ab^{-1} is a competitive facility in comparison with CEPC or Xenon1T.

The remaining parts of the paper are organized as follows. In Sec. II, we briefly discuss the direct and indirect detection constraints on the HSDM. In Sec. III we address the collider phenomenologies for the HSDM with dark matter mass in the narrow resonant mass region at the 14 TeV LHC and the 100 TeV FCC, where we focus on both the VBF channel and

mono- Z channel. Our main results are presented in Sec. IV, where we show the luminosities required for the 2σ exclusion and 5σ discovery. Finally we conclude in Sec. V.

II. MODEL AND CONSTRAINTS

A. Model

In the simplest HSDM model, the dark matter s communicates with the SM particles through the SM Higgs scalar. The Lagrangian for this mode reads as

$$\mathcal{L} = \mathcal{L}_{\text{SM}} + \frac{1}{2}(\partial s)^2 - \frac{\mu_s^2}{2}s^2 - \frac{\kappa_s}{2}s^2 |H|^2 - \frac{\lambda_s}{2}s^4, \quad (1)$$

where μ_s , λ_s and κ_s are the singlet scalar bare mass, the self-interaction coupling constant, and the coupling constant between dark matter and SM Higgs, respectively. A Z_2 parity, under which s is odd and other fields are even, is imposed to make the DM stable, which reduces the number of model parameters. After the electroweak symmetry breaking one can obtain

$$\mathcal{L} = \mathcal{L}_{\text{SM}} + \frac{1}{2}(\partial s)^2 - \frac{1}{2}m_s^2 s^2 - \frac{\kappa_s v}{2}s^2 h - \frac{\kappa_s}{4}s^2 h^2 - \frac{\lambda_s}{2}s^4, \quad (2)$$

where $m_s = \mu_s^2 + \kappa_s v^2/2$ is the physical mass of the singlet scalar, and $H = (v + h)/\sqrt{2}$, $s = \langle s \rangle + s$ and $v \simeq 246$ GeV.

Among the three model parameters, the self-interaction coupling λ_s does not directly affect the DM relic abundance, the DM-nucleon scattering cross section and DM production cross section at colliders, we simply decouple this parameter from the DM phenomenology discussed below. It turns out that the remaining two parameters are strongly constrained.

B. Constraints from indirect detections

Assume that the cold dark matter is saturated by the singlet scalar s , s should account for the DM relic density measured by the Planck and WMAP [46],

$$\Omega_{\text{DM}} h^2 = 0.1199 \pm 0.0027, \quad (3)$$

from which one infers the correlation between m_s and κ_s as shown in Fig.1. Besides the relic abundance in Eq.(3), there are other indirect constraints, including the Higgs invisible

decay $h \rightarrow ss$ in the mass region $m_s < m_h/2$ and the γ -ray limits from the Fermi-LAT [10, 11]. For the Higgs invisible decay, Fig.1 shows the latest limits at the 8 TeV LHC [12], HL-LHC and CEPC [44], which indicates that m_s below 52 GeV is excluded by the data $\text{Br}(h \rightarrow ss) \leq 29\%$, while the HL-LHC and CEPC can reach 54 GeV and 57 GeV, respectively.

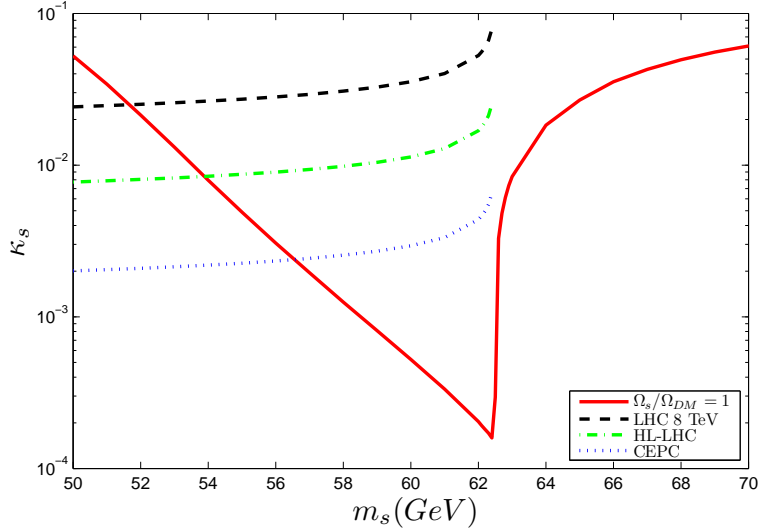


FIG. 1. Indirect constraints on the dark matter mass m_s from the dark matter relic abundance and Higgs invisible decays at the LHC Run 1, HL-LHC and CEPC.

C. Constraints from direct detections

The direct detection at LUX and Xenon1T can further constrain the parameter space, according to the spin-independent DM-nucleon scattering cross section given by

$$\sigma_{\text{SI}} = \frac{\kappa_s^2 f_N^2 \mu^2 m_N^2}{4\pi m_h^4 m_s^2}, \quad (4)$$

where m_N is the nucleon mass, $\mu = m_s m_N / (m_s + m_N)$ is the DM-nucleon reduced mass, and $f_N \sim 0.3$ is the hadron matrix element [28]. Fig.2 shows the predicted values of the spin-independent DM-nucleon scattering cross section, together with direct detection limits at XENON100 [8] and LUX [9] experiments. The limits at XENON1T [43] have been also shown. It indicates that the dark matter mass is restricted to a narrow resonant region between 53 GeV and 63 GeV. Once we employ the latest Fermi-LAT limits [33], this narrow mass region is further reduced to a narrow range between 53 GeV and 62.5 GeV.

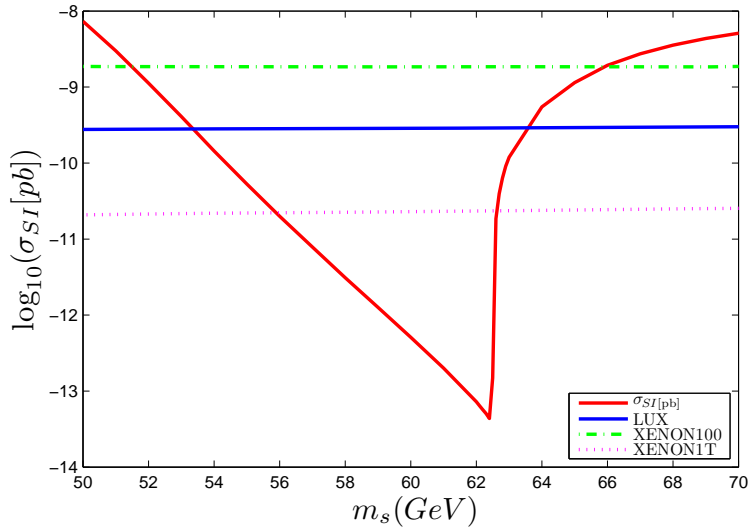


FIG. 2. Direct-detection constraints on dark matter mass m_s from LUX and Xenon experiments. The red curve represents the dark matter relic abundance constraint.

III. DARK MATTER AT HADRON COLLIDERS

In this section we study the collider signatures of the scalar dark matter at the 14 TeV LHC and 100 TeV FCC. We will explore the sensitivities at these two colliders for the dark matter mass in the narrow resonant region between 53 GeV and 62.5 GeV. We consider the dominant VBF channel as well as the sub-leading but relatively clean non- Z channel.

We use FeynRules [47] to generate model files prepared for MadGraph5 [48], which includes Pythia 6 [49] for parton showering and hadronization, and the package Delphes 3 [50] for fast detector simulation. In particular, the default CMS detector card and the Snowmass detector card are used for the 14 TeV and 100 TeV pp collider, respectively. Events are generated by using Madgraph5 at the leading order with the 125 GeV Higgs and fixed value $\kappa_s = 1.0$ for different dark matter masses. Cross sections are reproduced by rescaling κ_s^2 which corresponds to m_s .

A. Vector boson fusion

In the VBF channel, the dark matter pairs are produced through the Higgs scalar

$$pp \rightarrow h + jj \rightarrow ss + jj, \quad (5)$$

where the Higgs h should be on-shell in our case. The primary SM backgrounds to this process include Z +jets, W +jets, $t\bar{t}$ +jets and QCD multi-jets. For simplicity we consider the main contributions arising from Z + jets and W +jets channels, and adopt the cuts used by the CMS VBF analysis [51] for event selection:

$$\begin{aligned}
p_{T_{j_{1(2)}}} &> 50 \text{ GeV}, \quad |\eta_{j_{1(2)}}| < 4.7, \\
\eta_{j_1} \cdot \eta_{j_2} &< 0, \quad \Delta\eta_{jj} > 4.2, \\
M_{jj} &> 1100 \text{ GeV}, \quad \Delta\phi_{jj} < 1.0, \\
E_T^{\text{miss}} &> 130 \text{ GeV},
\end{aligned} \tag{6}$$

where $p_{T_{j_{1(2)}}}$ and $\eta_{j_{1(2)}}$ are the transverse momentum and pseudo-rapidity of the first (second) leading jet, respectively. $\Delta\eta_{jj}$, $\delta\phi_{jj}$ and M_{jj} are the rapidity difference, azimuthal-angle difference and invariant mass of the two leading jets, respectively. Any event with an additional jet with $p_T > 30$ GeV and pseudo-rapidity between those of the two tagged jets is rejected.

We firstly apply the criteria in Eq.(6) to the 8 TeV LHC with data of 19.5 fb^{-1} . The number of events for the SM background is 134 in the Z +jets channels and 145 in the W +jets channels, respectively. Compared with the reported number of events (99 in the Z + jets channels and 183 in the W + jets channels) by the CMS collaboration [51], they are consistent with each other.

The criteria in Eq.(6) will be also applied to both the 14 TeV LHC and 100 TeV FCC simulations for conservation. It is based on the following facts. First, there is little difference between the 8 TeV LHC and 14 TeV LHC except the collision energy, which means the cut on the pseudo-rapidity of the first two leading jets should remain unchanged. Second, the kinetic distribution of the signal events though on-shell Higgs decay and the main backgrounds Z +jets and W +jets have similar changing trends when one modifies the cuts in Eq.(6), as the mass difference between the Higgs and $Z(W)$ boson can be omitted compared to the variation of collision energy. It turns out that for a benchmark DM mass the effects on the ratio of signal over background events S/B are less than two times due to variations on the cuts in Eq.(6) such as $p_{T_{j_{1(2)}}} > \{40, 50, 60, 80, 100\} \text{ GeV}$, $\Delta\eta_{jj} > \{4, 4.2, 4.5\}$, $M_{jj} > \{900, 1100, 1300, 1500\} \text{ GeV}$, $\Delta\phi_{jj} < \{0.8, 1.0, 1.2\}$ and $E_T^{\text{miss}} > \{100, 120, 150, 180\} \text{ GeV}$. See table below for details.

$p_{T_{j1(2)}} > \{40, 50, 60, 80, 100\} \text{ GeV}$	$\Delta\eta_{jj} > \{4.0, 4.2, 4.5\}$	$M_{jj} > \{900, 1100, 1300, 1500\} \text{ GeV}$
$\{1, 1, 1.02, 0.97, 0.86\}$	$\{0.89, 1, 1.17\}$	$\{0.82, 1, 1.19, 1.38\}$
$\Delta\phi_{jj} < \{0.8, 1, 1.2\}$	$E_T^{\text{miss}} > \{100, 120, 130, 150, 180\} \text{ GeV}$	
$\{1.02, 1, 0.97\}$	$\{0.84, 0.95, 1, 1.09, 1.18\}$	

TABLE I. Effects on the ratio S/B due to variations on the cuts in Eq.(6) for benchmark DM mass $m_s = 53 \text{ GeV}$ at 100 TeV FCC.

B. Mono- Z channel

In the mono- Z channel the dark matter pairs are produced via the process

$$pp \rightarrow h + Z \rightarrow ss + Z. \quad (7)$$

Compared with the VBF channel, the mono- Z channel is sub-leading but relatively cleaner. For event selection in this channel we adopt the following cuts as suggested by the CMS leptonic mode analysis [52]:

$$\begin{aligned}
p_T^l &> 20 \text{ GeV}, \quad |\eta_{e(\mu)}| < 2.5(2.4), \\
|m_{ll} - m_Z| &< 10 \text{ GeV}, \quad E_T^{\text{miss}} > 80 \text{ GeV}, \\
p_T^{ll} &> 50 \text{ GeV}, \quad |u_{||}/p_T^{ll}| < 1.0, \\
\Delta\phi_{ll, \vec{p}_T^{\text{miss}}} &> 2.7, \quad |E_T^{\text{miss}} - p_T^{ll}|/p_T^{ll} < 0.2,
\end{aligned} \quad (8)$$

where p_T^{ll} is the dilepton transverse momentum and $u_{||}$ is defined as the component of $\vec{u} = -\vec{p}_T^{\text{miss}} - \vec{p}_T^{ll}$ parallel to the direction of \vec{p}_T^{ll} . Events are rejected if an additional electron or muon is reconstructed with $p_T > 10 \text{ GeV}$, and any event having two or more jets with $p_T > 30 \text{ GeV}$.

Similar to the discussions in the preceding section, the criteria in Eq.(8) are examined via the 8 TeV LHC simulation with data of 19.7 fb^{-1} . The number of events for the SM background is 88 in the mono- Z channel. Compared with the reported number of events 138 in the same channel by the CMS collaboration [52], they are consistent with each other. Unlike the CMS result, the next-leading order (NLO) effects have been neglected in our simulation. The two numbers of events likely match better if the NLO effects are included in our analysis. Following the similar facts mentioned in the preceding section, the criteria in Eq.(8) will be directly applied to both the 14 TeV LHC and 100 TeV FCC simulations.

IV. RESULTS

We present our main results in Fig.3 and Fig.4, which show the integrated luminosity \mathcal{L} needed for exclusion and discovery at the 14 TeV LHC and 100 TeV FCC, respectively. Here, we take the following definition about significance

$$\frac{S}{\sqrt{B}} \quad (\text{for exclusion}), \quad \frac{S}{\sqrt{S+B}} \quad (\text{for discovery}). \quad (9)$$

Systematic uncertainties are neglected in both the signal and the background simulations.

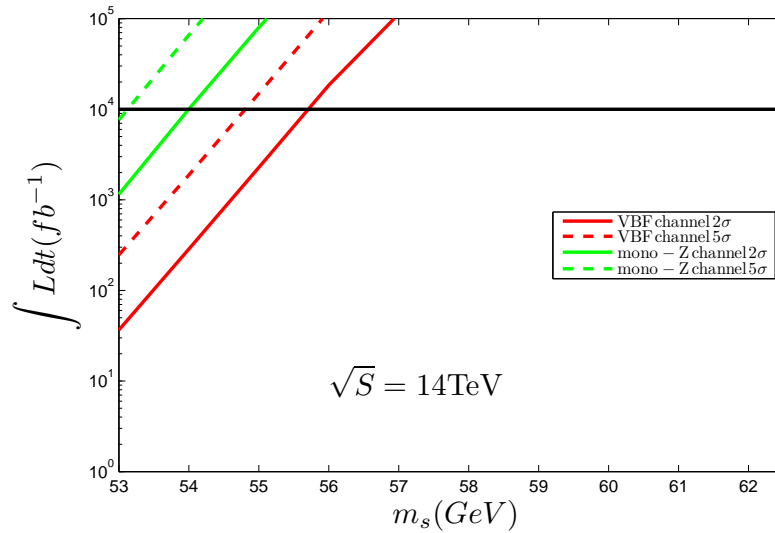


FIG. 3. The integrated luminosity needed for the exclusion determined by $S/\sqrt{B} = 1.96$ and 5σ discovery determined by $S/\sqrt{S+B} = 5$ in VBF and mono- Z channel at the 14 TeV LHC, respectively. The solid dark line represents the 10 ab^{-1} integrated luminosity.

In Fig.3 one observes that for $\mathcal{L} = 10^2 \text{ fb}^{-1}$ the scalar dark matter mass up to 53.5 GeV can be excluded, and for $\mathcal{L} = 10^3 \text{ fb}^{-1}$ the exclusion and discovery limits via VBF channel will reach to 54.6 GeV and 54 GeV, respectively. In contrast, Fig.4 shows that for $\mathcal{L} = 10^2 \text{ fb}^{-1}$ the scalar dark matter mass up to 54.5 GeV can be excluded, and for $\mathcal{L} = 10^3 \text{ fb}^{-1}$ the exclusion and discovery limits at the FCC via VBF channel approaches to 55.8 GeV and 55 GeV, respectively. The exclusion limits via the mono- Z channel are obviously weaker in comparison with the VBF channel.

Taking an integrated luminosity $\mathcal{L} = 10 \text{ ab}^{-1}$ at the 100 TeV FCC, one finds that the exclusion and discovery limits approach to 57 GeV and 56 GeV in the VBF channel, respectively. Compared with either CEPC or Xenon1T, where the exclusion limits approach

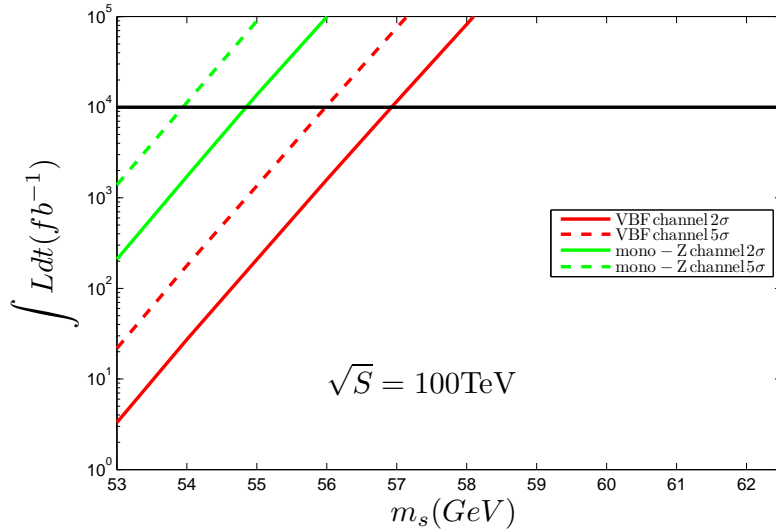


FIG. 4. Same as Fig.3, but for the 100 TeV FCC.

to 56.5 GeV and 56 GeV, respectively, the FCC with $\mathcal{L} = 10 \text{ ab}^{-1}$ is a competitive facility. Fig.3 and Fig.4 also illustrate that it is unlikely to detect the scalar dark matter in the mass range between 57 GeV and 62.5 GeV in HSDM model at any present and future facilities mentioned in this paper.

V. CONCLUSION

In this paper, we explored the collider signatures of the scalar dark matter in the HSDM model. Our study shows that for the 100 TeV FCC with an integrated luminosity of 10 ab^{-1} , the exclusion and discovery sensitivities reach to 57 GeV and 56 GeV respectively through the VBF channel, and 54.8 GeV and 53.9 GeV respectively via the mono-Z channel. Compared with either CEPC or Xenon1T, where the exclusion limits approach to 56.5 GeV and 56 GeV, respectively, FCC is a competitive facility. Unfortunately, the scalar dark matter in the mass range between 56.5 GeV and 62.5 GeV is unlikely to be either directly or indirectly detected at any present and future facility discussed in this paper.

ACKNOWLEDGMENTS

This work is supported in part by the National Natural Science Foundation of China under grant Nos. 11275245, 11135003 and 11405015, and by the CAS Center for Excellence

in Particle Physics (CCEPP).

- [1] G. Aad *et al.* [ATLAS Collaboration], Phys. Lett. B **716**, 1 (2012) [arXiv:1207.7214 [hep-ex]].
- [2] S. Chatrchyan *et al.* [CMS Collaboration], Phys. Lett. B **716**, 30 (2012) [arXiv:1207.7235 [hep-ex]].
- [3] V. Silveira and A. Zee, Phys. Lett. B **161**, 136 (1985).
- [4] J. McDonald, Phys. Rev. D **50**, 3637 (1994) [hep-ph/0702143].
- [5] J. McDonald, Phys. Rev. Lett. **88**, 091304 (2002) [hep-ph/0106249].
- [6] M. C. Bento, O. Bertolami, R. Rosenfeld and L. Teodoro, Phys. Rev. D **62**, 041302 (2000) [astro-ph/0003350].
- [7] C. P. Burgess, M. Pospelov and T. Veldhuis, Nucl. Phys. B **619**, 709 (2001) [hep-ph/0011335].
- [8] E. Aprile *et al.* [XENON100 Collaboration], Phys. Rev. Lett. **109**, 181301 (2012) [arXiv:1207.5988 [astro-ph.CO]].
- [9] D. S. Akerib *et al.* [LUX Collaboration], arXiv:1512.03506 [astro-ph.CO].
- [10] M. Ackermann *et al.* [Fermi-LAT Collaboration], Phys. Rev. Lett. **115**, 231301 (2015) [arXiv:1503.02641 [astro-ph.HE]].
- [11] M. Ackermann *et al.* [Fermi-LAT Collaboration], Phys. Rev. D **91**, 122002 (2015) [arXiv:1506.00013 [astro-ph.HE]].
- [12] G. Aad *et al.* [ATLAS Collaboration], ATLAS-CONF-2015-004.
- [13] H. Davoudiasl, R. Kitano, T. Li and H. Murayama, Phys. Lett. B **609**, 117 (2005) [hep-ph/0405097].
- [14] A. Kusenko, Phys. Rev. Lett. **97**, 241301 (2006) [hep-ph/0609081].
- [15] D. O'Connell, M. J. Ramsey-Musolf and M. B. Wise, Phys. Rev. D **75**, 037701 (2007) [hep-ph/0611014].
- [16] V. Barger, *et al.*, Phys. Rev. D **77**, 035005 (2008) arXiv:0706.4311 [hep-ph].
- [17] H. S. Cheon, S. K. Kang and C. S. Kim, JCAP **0805**, 004 (2008) [arXiv:0710.2416 [hep-ph]].
- [18] X. G. He, *et al.*, Phys. Rev. D **79** (2009) 023521 [arXiv:0811.0658 [hep-ph]].
- [19] R. N. Lerner and J. McDonald, Phys. Rev. D **80**, 123507 (2009) [arXiv:0909.0520 [hep-ph]].
- [20] M. Farina, D. Pappadopulo and A. Strumia, Phys. Lett. B **688** (2010) 329 [arXiv:0912.5038 [hep-ph]].

- [21] W. L. Guo and Y. L. Wu, JHEP **1010** (2010) 083 [arXiv:1006.2518 [hep-ph]].
- [22] S. Profumo, L. Ubaldi and C. Wainwright, Phys. Rev. D **82** (2010) 123514 [arXiv:1009.5377 [hep-ph]].
- [23] A. Biswas and D. Majumdar, Pramana **80**, 539 (2013) [arXiv:1102.3024 [hep-ph]].
- [24] A. Djouadi, O. Lebedev, Y. Mambrini and J. Quevillon, Phys. Lett. B **709**, 65 (2012) [arXiv:1112.3299 [hep-ph]].
- [25] C. Cheung, M. Papucci and K. M. Zurek, JHEP **1207**, 105 (2012) [arXiv:1203.5106 [hep-ph]].
- [26] A. Djouadi, A. Falkowski, Y. Mambrini and J. Quevillon, Eur. Phys. J. C **73** (2013) 2455 [arXiv:1205.3169 [hep-ph]].
- [27] J. M. Cline and K. Kainulainen, JCAP **1301**, 012 (2013) [arXiv:1210.4196 [hep-ph]].
- [28] J. M. Cline, K. Kainulainen, P. Scott and C. Weniger, Phys. Rev. D **88** (2013) 055025 [arXiv:1306.4710 [hep-ph]].
- [29] M. Endo and Y. Takaesu, Phys. Lett. B **743**, 228 (2015) [arXiv:1407.6882 [hep-ph]].
- [30] N. Craig, H. K. Lou, M. McCullough and A. Thalappilil, arXiv:1412.0258 [hep-ph].
- [31] C. E. Yaguna, JCAP **0903**, 003 (2009) [arXiv:0810.4267 [hep-ph]].
- [32] L. Feng, S. Profumo and L. Ubaldi, JHEP **1503**, 045 (2015) [arXiv:1412.1105 [hep-ph]].
- [33] M. Duerr, P. F. Perez and J. Smirnov, Phys. Lett. B **751**, 119 (2015) [arXiv:1508.04418 [hep-ph]].
- [34] M. Duerr, P. F. Perez and J. Smirnov, arXiv:1509.04282 [hep-ph].
- [35] M. Duerr, P. F. Perez and J. Smirnov, arXiv:1510.07562 [hep-ph].
- [36] H. Han and S. Zheng, JHEP **1512**, 044 (2015) [arXiv:1509.01765 [hep-ph]].
- [37] H. Han and S. Zheng, arXiv:1510.06165 [hep-ph].
- [38] N. Bernal and X. Chu, arXiv:1510.08527 [hep-ph].
- [39] E. Del Nobile, M. Nardecchia and P. Panci, arXiv:1512.05353 [hep-ph].
- [40] A. Beniwal, *et al.*, arXiv:1512.06458 [hep-ph].
- [41] Y. Mambrini, Phys. Rev. D **84**, 115017 (2011), arXiv:1108.0671 [hep-ph].
- [42] F. S. Queiroz, K. Sinha and A. Strumia, Phys. Rev. D **91**, no. 3, 035006 (2015), arXiv:1409.6301 [hep-ph].
- [43] E. Aprile *et al.* [XENON1T Collaboration], arXiv:1206.6288 [astro-ph.IM].
- [44] CEPC-SPPC Study Group Group (2015), http://cepc.ihep.ac.cn/preCDR/main_preCDR.pdf.

- [45] CEPC-SPPC Study Group Group (2015), http://cepc.ihep.ac.cn/preCDR/Pre-CDR_final_2015031.pdf.
- [46] P. A. R. Ade *et al.* [Planck Collaboration], *Astron. Astrophys.* **571**, A16 (2014) [arXiv:1303.5076 [astro-ph.CO]].
- [47] A. Alloul, *et al.*, *Comput. Phys. Commun.* **185**, 2250 (2014) [arXiv:1310.1921 [hep-ph]].
- [48] J. Alwall *et al.*, *JHEP* **1407**, 079 (2014) [arXiv:1405.0301 [hep-ph]].
- [49] T. Sjostrand, S. Mrenna and P. Z. Skands, *JHEP* **0605**, 026 (2006) [hep-ph/0603175].
- [50] J. de Favereau *et al.* [DELPHES 3 Collaboration], *JHEP* **1402**, 057 (2014) [arXiv:1307.6346 [hep-ex]].
- [51] S. Chatrchyan *et al.* [CMS Collaboration], *Eur. Phys. J. C* **74**, 2980 (2014) [arXiv:1404.1344 [hep-ex]].
- [52] V. Khachatryan *et al.* [CMS Collaboration], arXiv:1511.09375 [hep-ex].